



Special Section:

Jupiter Midway Through the Juno Mission

Key Points:

- Jupiter's asymmetric gravity field confines the vertical profile of the flow, with each harmonic constraining the flow at a different depth
- Deep zonal flows with a meridional profile significantly different from that of the cloud level are possible, but statistically unlikely
- A flow profile inspired by the microwave radiometer (MWR) measurements is consistent with the gravity measurements

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The Range of Jupiter's Flow Structures that Fit the Juno Asymmetric Gravity Measurements

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Abstract The asymmetric gravity field measured by the Juno spacecraft has allowed the estimation of the depth of Jupiter's zonal jets, showing that the winds extend approximately 3,000 km beneath the cloud level. This estimate was based on an analysis using a combination of all measured odd gravity harmonics, J_3 , J_5 , J_7 , and J_9 , but the wind profile's dependence on each of them separately has yet to be investigated. Furthermore, these calculations assumed the meridional profile of the cloud-level wind extends to depth. However, it is possible that the interior jet profile varies somewhat from that of the cloud level. Here we analyze in detail the possible meridional and vertical structure of Jupiter's deep jet streams that can match the gravity measurements. We find that each odd gravity harmonic constrains the flow at a different depth, with J_3 the most dominant at depths below 3,000 km, J_5 the most restrictive overall, whereas J_9 does not add any constraint on the flow if the other odd harmonics are considered. Interior flow profiles constructed from perturbations to the cloud-level winds allow a more extensive range of vertical wind profiles, yet when the meridional profiles differ substantially from the cloud level, the ability to match the gravity data significantly diminishes. Overall, we find that while interior wind profiles that do not resemble the cloud level are possible, they are statistically unlikely. Finally, inspired by the Juno microwave radiometer measurements, assuming the brightness temperature is dominated by the ammonia abundance, we find that depth-dependent flow profiles are still compatible with the gravity measurements.

Plain Language Summary Jupiter's north-south asymmetric gravity field, as measured by the Juno spacecraft, currently orbiting Jupiter, has been used to determine that the depth of its jet streams (associated with the famous visible cloud bands) are approximately 3,000 km. This estimate was based on all the gravity field measurements combined. However, there is also information about the structure of the flow hidden in each individual measurement. Here we analyze these measurements and show how each of them is sensitive to the flow at a different depth. We also systematically investigate the statistical likelihood of deep wind profiles that differ from the profile observed at the cloud level with various structures at depth. We find that Jupiter's measured cloud-level jet streams fit with its gravity data for a relatively narrow envelope of vertical structures. Although other jet profiles that are different from the one observed at the cloud level are feasible (still consistent with the gravity data), they are statistically unlikely. Finally, we explore a depth-dependent wind structure inspired by the Juno microwave radiometer instrument, which indicates that ammonia abundance varies with depth and might be correlated with the jet streams. We find that such a profile can still match the gravity data as long as the variation from the cloud-level wind is not substantial.

1. Introduction

The Juno spacecraft has provided an unprecedented glance into Jupiter's atmospheric flows below the cloud level. The high-precision gravity measurements, particularly those of the odd gravitational harmonics, repeated in multiple passes (Iess et al., 2018), have presented an opportunity to estimate the depth and structure of Jupiter's zonal jets. It has been found that the zonal jets are deep and penetrate to approximately 3,000 km below the cloud level (Kaspi et al., 2018). Below this depth, the even gravitational harmonics indicate that Jupiter rotates almost like a solid body (Guillot et al., 2018). However, determining the details of the decay profile with depth poses a significant challenge. Remnants of the zonal flows appear even below 4,000 km, and since the estimation of the electrical conductivity in Jupiter at this depth is at least 10 S m^{-1} (Nellis et al., 1996; Wicht, Gastine, & Duarte, 2019; Wicht, Gastine, Duarte, & Dietrich, 2019), an interaction between the flow and the magnetic field is expected there (Cao & Stevenson, 2017; Duer et al., 2019; Galanti et al., 2017; Moore et al., 2019). Understanding the gravity harmonic signature and the flow structure below the cloud level is thus essential in order to build a better picture of Jupiter's atmosphere.

The gravity field of Jupiter, represented by the gravity harmonics, reflects both the internal density structure and the deep zonal flow structure (Hubbard, 1999; Kaspi et al., 2010). The even gravity harmonics are used to constrain the internal density structures of Jupiter and other gas giants (e.g., Helled et al., 2010; Hubbard et al., 1974, 1975; Nettelmann et al., 2013). Multiple studies have shown that the higher-order (even) gravity harmonics are sensitive to the outer regions of the planet (e.g., Guillot & Gautier, 2007; Nettelmann et al., 2013; Zharkov & Trubitsyn, 1974). Their exact value is defined by the density distribution throughout the planet and the planet's rotation, composition, shape, mass, and radius. Since for a static gas planet, the odd harmonics are identically zero, any gravitational asymmetry between north and south would indicate a dynamical source generating those asymmetries (Kaspi, 2013). Juno measured with high precision the gravity harmonics up to J_{10} , including significant odd values. The measured values and error range are $J_3 = (-4.24 \pm 0.91) \times 10^{-8}$, $J_5 = (-6.89 \pm 0.81) \times 10^{-8}$, $J_7 = (12.39 \pm 1.68) \times 10^{-8}$, and $J_9 = (-10.58 \pm 4.35) \times 10^{-8}$ (Iess et al., 2018). The relation between the density anomaly and the flow (thermal wind balance) allows constraining the deep flow structure within the planet (Kaspi, 2013; Kaspi et al., 2010, 2018). Assuming that the cloud-level zonal wind profile is extended towards Jupiter's interior using a scaling factor, one can find many solutions for the deep flow structure that satisfy all four odd gravity harmonics within the uncertainty range. With the currently available data, Jupiter's deep flow cannot be determined uniquely (Kaspi et al., 2018; Kong et al., 2018), and systematic exploration of the range of the deep flow structure is necessary.

Moreover, the meridional profile of the zonal wind is not necessarily constant with depth. The cloud-level wind itself has a measurement error (Garcia-Melendo & Sánchez-Lavega, 2001; Tollefson et al., 2017; Salyk et al., 2006), and as it extends inward, the profile might vary, although any such variation must be accompanied with a meridional temperature gradient as well. Some evidence for such meridional variations come from the Juno microwave radiometer (MWR) measurements showing that the nadir brightness temperature profile (dominated by the ammonia abundance) becomes smoother with depth (Bolton et al., 2017; Li et al., 2017). Although this measurement does not necessarily correlate with temperature, it does coincide, to some degree, with the zonal wind profile at the cloud level (Bolton et al., 2017) and thus might hint to the vertical variation of the zonal wind profile in the upper 300 km.

Previous work on constraining the deep flow structure was done using all four measured gravity harmonics combined (e.g., Kaspi et al., 2018, 2020; Kong et al., 2018). However, an important question is how does each gravity harmonic individually constrain the flow strength at different depths. Here, we examine the individual contribution of each odd gravity harmonic, with emphasis on the depth of influence and the relation to the cloud-level zonal wind profile. In order to provide a systematic analysis, we take a hierarchical approach, in which we increase the level of complexity of the variation of the wind structure and in all cases explore what is the range of solutions that match the gravity measurements. We begin with solutions that are identical to the cloud-level profile and allow only for the vertical decay to vary. Then, we relax the constraint on the meridional profile of the zonal wind and allow variations from the measured cloud-level profile along with varying the vertical decay. Finally, we examine random meridional profiles that are not related at all to Jupiter's measured cloud-level profile and explore the possibility that the interior wind structure, which influences the gravity measurements, is completely different from the cloud-level flow. Following this logic, we also search for solutions with smoother wind profiles that resemble the MWR measurements at 300 km (Channel 1) and calculate the vertical profile of such flows that can match also the gravity data.

The paper is organized as follows: In section 2, we introduce the theoretical background for this analysis, connecting the gravity measurements and the wind profile. In section 3, we present the possible solutions for Jupiter's wind profile, the depth sensitivity obtained by excluding a specific harmonic, and the contribution function of each harmonic. In section 4, we discuss the ability to find solutions for the anomalous gravity field of different meridional profiles, and in section 5, we explore depth-dependent meridional structures, inspired by the MWR measurements. We discuss the dynamical implications of the results and conclude in section 6.

2. Methodology

The density distribution within Jupiter is reflected in the zonal gravity harmonics (J_n), which describe the external gravitational field of the planet in equilibrium (Zharkov & Trubitsyn, 1974). The gravity harmonics can be represented by

$$J_n = -\frac{1}{MR_J^n} \int \rho r^n P_n(\mu) d^3 r, \quad (1)$$

where M and R_J are Jupiter's mass and equatorial radius, respectively, n is the harmonic degree ($n=2, \dots, N$), ρ is the density, r is the radial coordinate, and $P_n(\mu)$ is the n th Legendre polynomial, where $\mu = \sin\theta$ and θ is the latitude (Hubbard, 1984). The density can be decomposed such that $\rho(r, \theta) = \tilde{\rho}(r, \theta) + \rho'(r, \theta)$, where $\tilde{\rho}(r, \theta)$ is the static component that is determined by the planet's shape and rotation (Hubbard, 2012) and $\rho'(r, \theta)$ is the dynamical anomaly representing fluid velocities with respect to the solid body rotation (Kaspi et al., 2010). The zonal gravity harmonics that represent only the dynamical part of the flow (ΔJ_n) can be calculated by integrating the density anomaly and its projection onto the Legendre polynomials in spherical coordinates such that

$$\Delta J_n = -\frac{2\pi}{MR_J^n} \int_0^{R_J} \int_{-1}^1 \rho'(r, \mu) r^{n+2} P_n(\mu) d\mu dr. \quad (2)$$

Since an oblate planet with no dynamics is symmetric between north and south, the density anomaly represented by the odd harmonics ($n=3, 5, \dots$) should be identically zero if the flow pattern is symmetric and will be very small if the dynamics are shallow ($\Delta J_n = J_n$ for odd n). However, Juno measured four odd gravity harmonics (Iess et al., 2018), indicating the existence of a strong asymmetric pattern in Jupiter's flow field due to the existence of strong, deep winds.

The rapid rotation and size of the planet (small Rossby number) imply that this asymmetry is directly related to zonal flows, since, to first order, the leading balance in Jupiter is a geostrophic balance between the flow-related Coriolis forces and the pressure gradients. This leads to a vorticity balance known as thermal wind balance (Kaspi et al., 2009; Pedlosky, 1987). If only zonal (azimuthal) flows are considered, the thermal wind balance can be written as

$$2\Omega r \frac{\partial(\tilde{\rho}u)}{\partial z} = g_0 \frac{\partial \rho'}{\partial \theta}, \quad (3)$$

where Ω is Jupiter's rotation rate, $u(r, \theta)$ is the zonal flow, $g_0(r)$ is the mean gravitational acceleration, and z is the direction parallel to the rotation axis. An equivalent equation can be written with temperature instead of density gradients, and one can easily switch between the two versions through the equation of state. Note that the barotropic limit of Equation 3 is not simply when the right-hand side vanishes, but when $\frac{\partial u}{\partial z} = 0$ (see full derivation in Kaspi et al., 2016). Galanti et al. (2017) showed that a higher-order expansion, beyond thermal wind balance, only slightly adjusts the deep flow dynamics (by less than 10%). Therefore, for the purpose of studying the overall vertical profile, Equation 3 is a good approximation.

Our goal here is to search for possible deep wind structures that can explain each of the measured odd gravity harmonics (J_3, J_5, J_7 , and J_9). Unlike previous studies (e.g., Kaspi et al., 2018), we are not solving for an optimal solution with respect to the full error covariance matrix. Any vertical wind profile that fits the odd measured gravity harmonics, within the uncertainty range of Juno, is considered a possible solution for the flow. This allows us to examine the full range of possible solutions, without converging on a single decay profile of the flow. For example, the optimal solution suggested by Kaspi et al. (2018) that considered the error covariance matrix is not a solution here since the value of J_3 is not within the measured error.

3. The Vertical Profile of the Zonal Flow

Taking a hierachal approach entailing an increasing level of complexity, we first use the observed cloud-level wind as an upper boundary condition for the flow field and assume the same profile continues inward in a direction parallel to the spin axis, due to angular momentum considerations (Busse, 1976; Kaspi et al., 2010). The possible deep flow structures are then set to decay continuously from the cloud level to a few thousand kilometers below it (Kaspi et al., 2018), using two different decay regions. Dividing the

decay functions into two distinct regions stems from the magnetic field's possible effects on the flow, expected approximately at $r < 0.97 R_J$ (Duer et al., 2019; Wicht, Gastine, Duarte, & Dietrich, 2019), which imply that once the electrical conductivity becomes dominant, the magnetic field acts to dissipate the flow (Gastine et al., 2014; Liu et al., 2008). Thus, for the lower part (the semiconducting region), we chose an exponential decay (Equation 5, $r < R_T$) that fits the exponential nature of the electrical conductivity within Jupiter (French et al., 2012; Nellis et al., 1992; Weir et al., 1996). For the upper part, the vertical decay function includes both an exponent and hyperbolic tangent (Equation 5, $R_T \leq r \leq R_J$), which combine to give a wide range of possible decay profiles.

The vertical profile of the zonal flow is defined with six independent parameters, chosen to cover an extensive range of vertical profiles. It is set as

$$u(\theta, r) = u_{\text{proj}}(\theta, r)Q_s(r), \quad (4)$$

$$Q_s(r) = \begin{cases} (1 - \alpha)\exp\left(\frac{r - R_J}{H_1}\right) + \alpha \left[\frac{\tanh\left(\frac{-R_J - H_2 - r}{\Delta H}\right) + 1}{\tanh\left(\frac{H_2}{\Delta H}\right) + 1} \right] & R_T \leq r \leq R_J \\ Q_s(R_T)\exp\left(\frac{r - R_T}{H_3}\right) & r < R_T, \end{cases} \quad (5)$$

where $u_{\text{proj}}(r, \theta)$ is the wind at the cloud level, projected inward (with no decay) in the direction parallel to the axis of rotation (\hat{z} axis, Equation 3), $Q_s(r)$ is the radial decay function, representing the fraction of the cloud-level wind at every depth, and the set of parameters that forms the decay are bounded by the following limits: $0 \leq \alpha \leq 1$, $200 \leq H_1 \leq 2500$ km, $200 \leq H_2 \leq 2500$ km, $200 \leq \Delta H \leq 2500$ km, $0.95 \leq R_T \leq 0.975 R_J$, and $100 \leq H_3 \leq 900$ km. The function Q_s is also smoothed at the transition depth. We vary the parameters uniformly between their lower and upper bounds, taking only profiles where the wind speed decays monotonically as viable options. In total, we consider 5×10^5 decay profiles, sufficiently covering the parameter space. This set of decay profiles serves as the sample population for this study. For each profile, we calculate the associated density anomaly and the implied odd gravity harmonics. All the calculations presented below are performed using the same sample population. Note that other forms of Q_s are possible and can still fit the measured gravity data ($Q_s(r, \theta)$, for example, as explored in Kaspi et al., 2018). However, we have found that for the exploration of the individual gravity harmonic depth sensitivities and the meridional profile anomalies, the chosen function, which allows a very wide range of decay profiles, is satisfactory.

From the 5×10^5 decay options examined, 6,712 vertical profiles are compatible with Juno's measured odd gravity harmonics, which represent a little over 1% of the sample population (Figure 1a). All the compatible decay profiles are located in a relatively narrow envelope, especially in the region around 2,000 km depth and the one below 4,000 km, with all the options pointing to remnants of jet-associated velocities at a depth of 4,000 km (Figure 1). Those deep velocities are still of the order of 1 m s^{-1} and despite being small, they are still higher than the magnetic secular variation associated velocities estimated by Moore et al. (2019). Increasing the error range of Juno's gravity measurements does allow for more solutions, but the overall structure does not change much (Figure 1b).

3.1. The Depth Sensitivity of the Odd Harmonics

Research to date has focused on finding vertical profiles that match all four odd gravity harmonics. However, there is information to be obtained from each gravity harmonic separately. Here, vertical flow profiles that fit three out of the four measured odd gravity harmonics are considered, and the depth sensitivity of the excluded harmonic is studied by examining the difference between the vertical profiles that include the specific J_n to those that do not necessarily include it. The resulting depth sensitivity of each odd measured gravity harmonic, according to Jupiter's measured meridional profile of the zonal wind, is presented in Figure 2. The gray envelope, the same one as in Figure 1b, is the boundary of all the possible solutions that fit all four odd gravity harmonics within 3σ . Note that not all possible profiles inside the gray envelope will necessarily

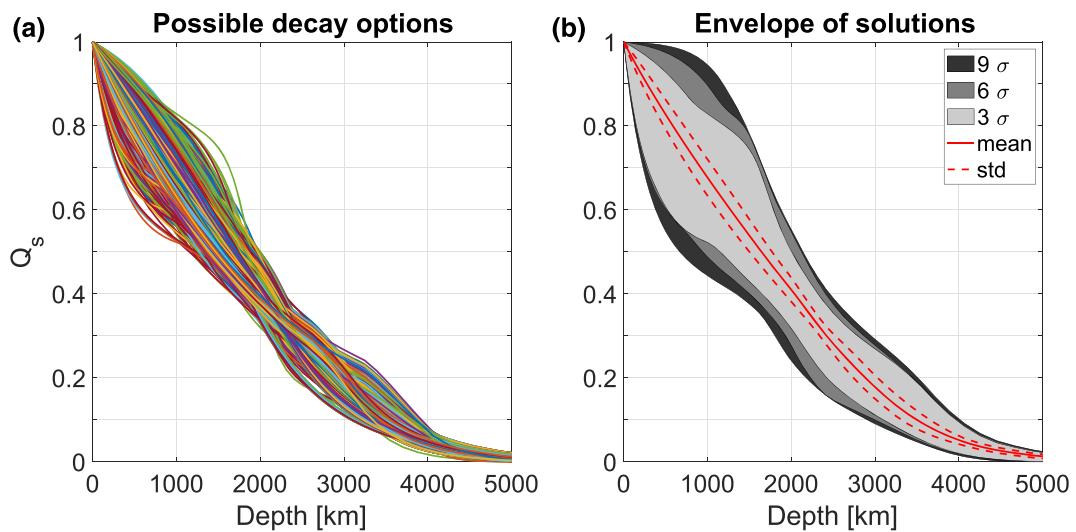


Figure 1. (a) Decay options that fit all four measured odd gravity harmonics (J_n) within the 3σ uncertainty of Juno, using Jupiter's observed cloud-level flows (Tollefson et al., 2017). (b) The envelope of all possible solutions (light gray), the average solution of all options (red line), their standard deviation (dashed red lines), and an envelope of solutions that satisfy a larger uncertainty range (darker gray with rising uncertainty range).

generate a solution compatible with the measured gravity field, since the solution is also dependent on the decay profile within the given envelope. All the other solutions gained while excluding one of the odd gravity harmonics appear in Figure 2 (turquoise envelopes). The turquoise envelopes always contain the gray envelopes by definition, since they are constructed by fitting at least three gravity harmonics. The difference between the turquoise envelopes and the gray ones denotes the region in which the *excluded* harmonic bounds the flow.

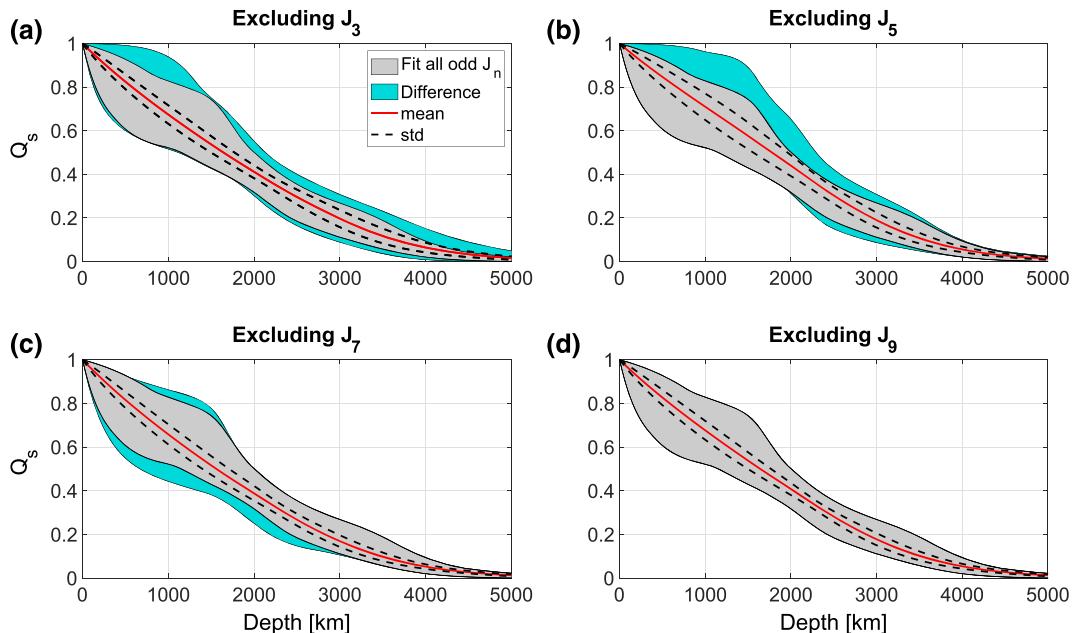


Figure 2. (a) The envelope of possible solutions that fit all four odd gravity harmonics (gray), the envelope of additional solutions once excluding J_3 (while still fitting J_5 , J_7 , and J_9) (turquoise), the average of all decay options within the panel (gray and turquoise combined) (red), and their standard deviation (dashed black). The other panels are the same while excluding other J_n , (b) excluding J_5 , (c) excluding J_7 , and (d) excluding J_9 . The results are a combination of the sample decay options (5×10^5). Note that the turquoise color emphasizes the depth sensitivity of each gravity harmonic separately.

The most insignificant influence is clearly of J_9 (Figure 2d). It appears to add no solutions at all to the gray envelope, meaning that J_9 does not constrain the flow if the other three odd values are still within Juno's 3σ . This is likely because J_9 has the highest measurement error and lowest signal-to-noise ratio (SNR), so even while fitting J_9 , there is an extensive region of solutions, and excluding it does not add new solutions. The largest influence on the flow profile and depth sensitivity comes from J_5 (Figure 2b). It appears to set the upper boundary of the gray envelope from the cloud level (0 km) to 3,500 km and a lower boundary of the gray envelope between 2,000 and 3,500 km. The strongest sensitivity is between the cloud level and 3,000 km. J_5 has the smallest measured 3σ value and largest SNR, but its value is very similar to the SNR of J_7 , so the large influence of J_5 cannot be a result of the SNR alone. In a similar manner, J_3 is mostly sensitive between 3,000 and 5,000 km and between the cloud level (0 km) and 1,500 km (Figure 2a). Note that a flow profile that decays to zero at 4,000 km ($\sim 0.94 R_J$) cannot fit J_3 . J_7 is sensitive between 500 and 2,500 km and sets mainly the lower boundary of the gray envelope at those depths (Figure 2c).

Previous studies that examined the depth dependency of the even gravity harmonics (resulting from the shape and density of a solid-body model, without differential flows) concluded that higher-order harmonics are more sensitive to the density in the outer regions (e.g., Guillot & Gautier, 2007; Nettelmann et al., 2013; Zharkov & Trubitsyn, 1974). This is implied by the radial dependence of the gravity harmonics (Equation 1). However, the above analysis shows that the wind-induced odd harmonics' depth dependency is more complicated. One exception is J_3 , the only harmonic that is sensitive below 4,000 km, which resembles the even harmonics' depth tendency, where the low-order harmonics are more sensitive in deeper regions.

3.2. The Contribution Function

The depth sensitivity of the gravity harmonics can also be examined by calculating directly the depth dependence of J_n , defined as the contribution function. This function was calculated in past studies for the even harmonics of Jupiter and other planets (e.g., Guillot & Gautier, 2007; Helled et al., 2010; Nettelmann et al., 2013). The contribution of each shell is the normalized integrant of J_n , defined as

$$C_n = \frac{1}{J_n} \frac{dJ_n}{dr} = \frac{1}{J_n} \frac{-2\pi}{MR_J^n} \int_{-1}^1 \rho(r, \mu) r^n + {}^2P_n(\mu) d\mu \quad (6)$$

(Hubbard, 1984; Hubbard et al., 1974; Zharkov & Trubitsyn, 1974). The even harmonics in past studies were calculated from the background density (solid body models), while in our study we use the wind-induced anomalous density field to calculate the odd harmonics' contribution, taking ρ' instead of ρ in Equation 6.

The averaged anomalous density profile of all possible decay structures, that are consistent with the four odd gravity harmonics, is presented in Figure 3a. The anomalous density reveals a change of sign at 2,000 km. The averaged odd contribution functions (C_n) and their standard deviations (Figure 3b), corresponding to the solution envelope from Figure 1, show a consistent sign change. Note that the change of sign is exhibited only by the anomalous density, corresponding to the wind shear with depth, and, therefore, does not exist when examining the even harmonics resulting from the static density (e.g., Nettelmann et al., 2013). The integrals of the nonnormalized contribution curve, C_n , are the gravity harmonic values, J_n , so the sign and value of J_n are set by the difference between the positive and negative curves (above and below 2,000 km, not shown). For the averaged anomalous density, the gravity harmonics are: $J_3 = -4.29 \times 10^{-8}$, $J_5 = -7.50 \times 10^{-8}$, $J_7 = 10.8 \times 10^{-8}$, and $J_9 = -6.69 \times 10^{-8}$.

The contribution function reveals a complex depth dependence for all four gravity harmonics. The depth sensitivity of each contribution function is marked by triangles (Figure 3b), which represent the depth of the mean absolute anomaly. The contribution function of J_3 , C_3 , has the largest areas under the curves at both the shallower (0–2,000 km) and deeper regions (>2,000 km). The depth of the mean anomaly, which here equals 2,020 km (Figure 3b, blue triangle), is near the depth of the sign change (2,000 km), meaning that J_3 gets near-equal anomalies from both regions. The standard deviation of C_3 (Figure 3b, blue shading) is the largest, implying a large variability of the solutions with depth when considering the J_3 value. The mean anomaly of C_5 is located in the deeper part of the domain (Figure 3b, red triangle), and the standard

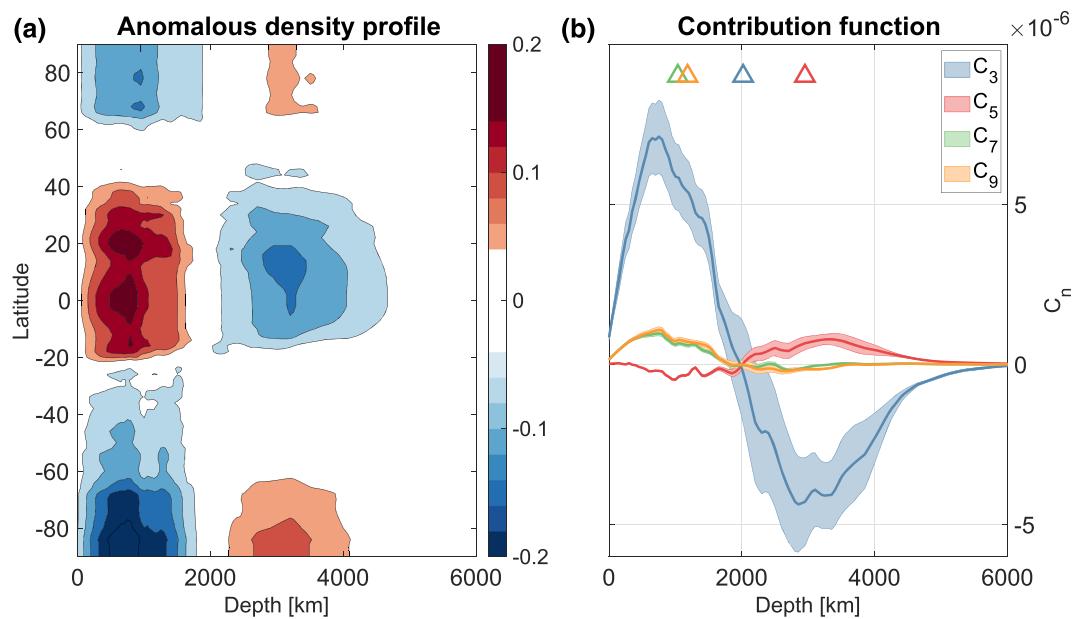


Figure 3. (a) The mean anomalous density profile of all possible decay options that fits the Juno four measured odd gravity harmonics; colors represent the anomalous density values (kg m^{-3}). (b) Averaged contribution function (lines) (m^{-1}) for each of the odd gravity harmonics and their associated standard deviation (shading); triangles represent the depth of the mean anomaly. Both panels are for all the latitudes and for only the upper $\sim 6,000$ km, below which the anomalous density is 0.

deviation of C_5 is substantial only between 2,000 and 4,000 km. C_5 is the only harmonic dominated mostly by the deeper region, emphasizing the important effect of J_5 on the deep wind structure (section 3.1). The mean anomalies of C_7 and C_9 are clearly located in the shallower region (yellow and green triangles $< 2,000$ km), and their standard deviation is small everywhere. The contribution of both C_7 and C_9 is 0 below 3,000 km, corresponding to Figure 2. Since C_7 and C_9 lay almost on top of each other, C_7 might mask the depth dependency of C_9 , as revealed in Figure 2 (along with the low SNR of J_9), so that if J_7 is within Juno's error range, so is J_9 . It is evident that the contribution function of the odd harmonics exhibits a more complicated pattern than the classical even harmonics (e.g., Guillot & Gautier, 2007; Helled et al., 2010; Nettelmann et al., 2013). As in the previous analysis, we find that the higher-order odd harmonics are not simply more pronounced in the outer regions. The projection of the wind patterns onto different depths is reflected in the odd harmonics' contribution at those depths, suppressing the $(r/R_J)^n$ dependency, which is the prominent feature of the even harmonics' contribution.

4. Sensitivity to the Meridional Profile of the Zonal Flows

Next, we relax the assumption, used in the previous section, that the meridional profile of Jupiter's zonal flow remains constant at all depths. First, the zonal wind profile is measured by tracking cloud motion, which itself has some uncertainty (Tollefson et al., 2017). Second, and most importantly, the assumption that the cloud-level profile extends perfectly to depth along the direction of the spin axis requires the flow to be locally close to barotropic (in the upper few thousand kilometers), which is not necessarily the case. Although the flow cannot be completely barotropic if $Q_s \neq 1$ (Equation 4), the horizontal density gradients required to balance the vertical changes associated with Q_s may be small (Equation 3). Any further deviation from close-to-barotropic flow must be supported by horizontal density (or temperature) gradients, which themselves must be maintained by some internal mechanism (Showman & Kaspi, 2013). Internal convection models support the scenario that there may be internal shear over the upper few thousand kilometers, but the overall structure of the flow does not change much (Jones & Kuzanyan, 2009; Kaspi et al., 2009). Any significant deviation from the zonal wind profile observed at the cloud level requires significant shear and, therefore, notable horizontal thermal gradients (thermal-wind balance). As this is an open question, for the

purpose of this analysis, we examine several cases of zonal wind meridional profiles, under the assumption that the wind profile possibly varies close to the cloud level and then projects inward without further modifications. For the purpose of the gravity analysis, this means that the altered meridional profiles occupy enough mass to affect the gravity field, and the flow observed at the cloud level is limited to a shallow-enough layer so it does not affect the gravity field.

The simplest case is clearly to use the measured profile at Jupiter's cloud level and allow its magnitude to decay with depth (section 3). A slightly less constraining option is to insert a perturbation into the measured profile, thereby keeping the general form and allowing a varying level of modifications to the cloud-level flow. The perturbed winds chosen here might represent the measured uncertainties in Jupiter's cloud-level wind (Garcia-Melendo & Sánchez-Lavega, 2001; Tollefson et al., 2017). Finally, random meridional profiles of the zonal flow with a spectra generally similar to that of Jupiter are examined as well.

The modified zonal flow profile is chosen at the cloud level and projected inward along the rotation axis (u_{proj} , Equation 4) with a range of vertical profiles, as described in section 3. The profiles are calculated by adding sinusoidal perturbations to the measured wind. The modified profiles have a standard deviation of $5 \pm 0.5 \text{ ms}^{-1}$ (varies with latitude) relative to the cloud-level winds, well within the measurement error (Garcia-Melendo & Sánchez-Lavega, 2001; Tollefson et al., 2017). The perturbation is constructed as

$$\epsilon(\theta) = \sum_{n=1}^{10} [a_n \sin(2n\theta) + b_n \cos(2n\theta)], \quad (7)$$

where ϵ is the perturbation and a_n and b_n are random numbers that are normally distributed around 0 with a standard deviation of 2 m s^{-1} . We first examine 1,000 modified profiles, each constructed by adding the perturbation to the measured wind (section 4.1). In addition, 1,000 random profiles are constructed purely from the ϵ function (Equation 7), where a_n and b_n have a standard deviation of 30 m s^{-1} . These profiles represent internal winds that are completely unrelated to the observed cloud-level winds (section 4.2).

4.1. Perturbed Cloud-Level Wind Profiles

The perturbed wind profiles (Figure 4a, colors) result in a substantially bigger solution envelope (Figures 4b.1–4b.4, gray) than the one from the measured zonal wind profile case, consistent with the fact that a wider range of wind profiles is allowed. Note that the overall shape has changed and that the flow can even vanish at $\sim 2,500 \text{ km}$. This might have an important implication, since the initial time-dependent magnetic field results from Juno imply that the wind at these depths are very weak (Duer et al., 2019; Moore et al., 2019). An important result is that even for the perturbed winds there are no solutions that fit at least three odd J_n that vanish above 2,000 km. The depth sensitivity of each harmonic is less pronounced than for the measured wind case. This reflects the fact that Figure 4 is a combination of all the possible solutions from 1,000 examined meridional structures. Overall, J_3 is still sensitive in the deeper regions (exemplified by the mean profile being weaker at depth, red line Figure 4b.1), although J_7 and J_9 contribute at depth as well. J_5 turns out to be the most insignificant harmonic and J_9 does affect the depth range of 1,500–2,000 km, unlike in the unperturbed wind case. The substantially larger range of solutions, however, does not manifest in more solutions relative to the examined cases. From 1,000 examined profiles individually tested with the decay sample population, only about 0.1% fit the anomalous gravity field compared to about 1% in the unperturbed case (Figure 5, red and blue). This suggests that although perturbed cloud-level wind profiles are possible, it is statistically more likely that a profile that is similar to the projection of the observed cloud-level wind is indeed the profile in the deeper atmosphere of Jupiter.

4.2. The Possibility of Other Zonal Wind Profiles

Next, we consider profiles that do not resemble Jupiter's cloud-level winds (Figure 6a). The resulting solution envelopes of the other zonal wind profiles are relatively similar to the previous case of perturbed winds (not shown). Only a very small subset of profiles (13 meridional profiles out of 1,000, about 1%) fit the four measured odd gravity harmonics (Figure 6b.1). The possibility of fitting two or more odd harmonics is rare and exists in only 7% or less of the zonal wind meridional profiles examined (Figures 6b.2 and 6b.3). J_3 is the harmonic that is pronounced in the majority of profiles (Figure 6b.4). In 14% of the examined random profiles, no odd harmonic is within the sensitivity range. In general, the measured harmonic's alignment with the

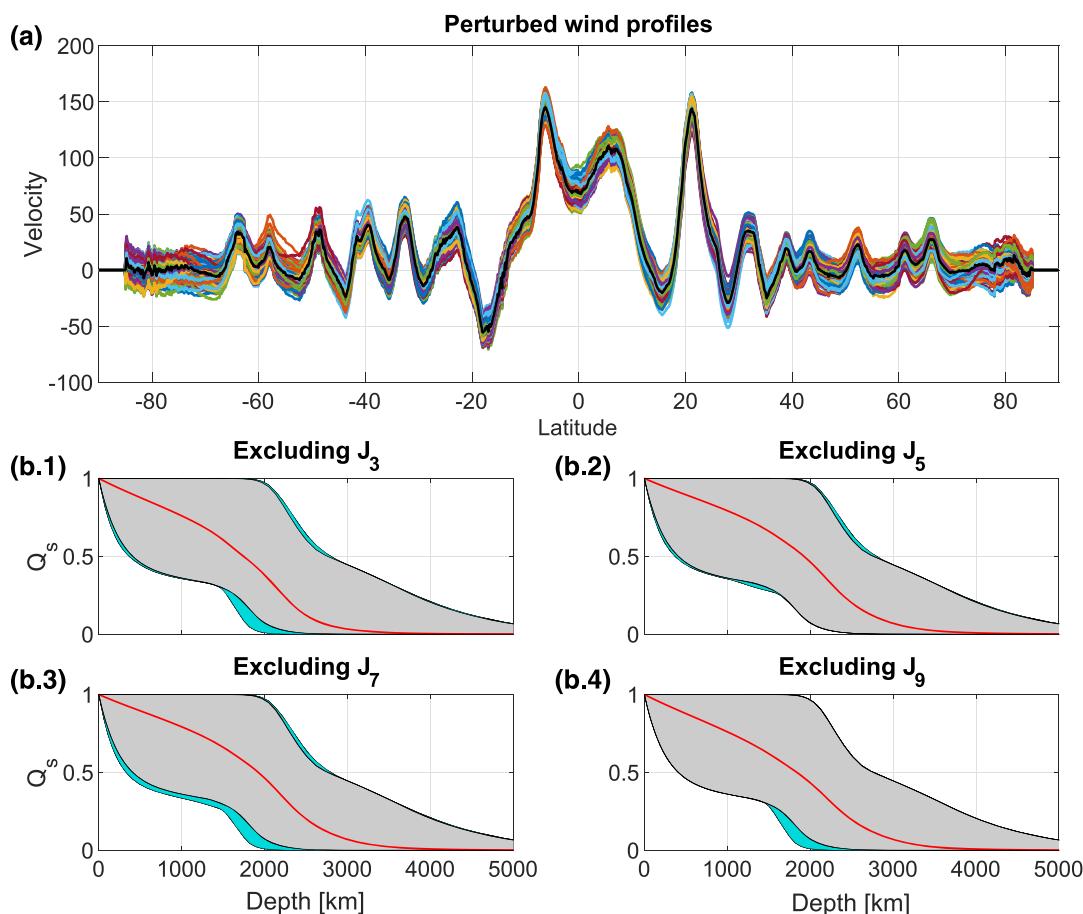


Figure 4. (a) One hundred examples of the 1,000 perturbed wind profiles (colors, [m s⁻¹]) and Jupiter's measured wind profile (black). (b.1–b.4) The odd gravity harmonics of the perturbed wind profiles depth sensitivity summary as in Figure 2. Each profile is examined with the same set of decay options. The results shown here are for all meridional and vertical zonal flow profiles options combined.

zonal flow structure does not appear to be coincidental. These findings are expected, considering that it is unlikely that an utterly different meridional profiles arise below the cloud level of Jupiter.

The ability of the 1,000 examined random profiles, each with its sample of decay options, to fit all four odd gravity harmonics is considerably smaller than previous cases—only about 0.01% (Figure 5, orange). This indicates that fitting all four odd harmonics is difficult with random meridional profiles of zonal wind. A summary of the examined cases appears in Figure 5. Note that the ordinate is a logarithmic scale and that 100% stands for all the meridional profiles (1,000 zonal wind profiles other than the measured cloud-level wind) and all decay options (5×10^5) for each case. We find that the envelope of possible solutions from Figure 1 stands for $\sim 1\%$ of the tested vertical profiles for zonal flows. The fitting percentage decreases with increasing perturbations, and drops rapidly when switching to random profiles. This trend repeats for all variations of at least three odd harmonics. For all cases, the random winds show a significantly lower fitting percentage than the other cases. We further present the fitting percentage obtained following the exclusion of two and three harmonics. In summary, we find that other meridional profiles of the zonal wind are possible, but they are statistically unlikely. This result implies that the meridional profile of Jupiter's zonal winds extends into the interior along the direction of the spin axis and weakens with depth, and is likely not significantly different from the cloud-level profile.

5. Zonal Wind Profiles Inspired by the MWR Measurements

In addition to the gravity measurements, Juno's six-channel MWR measurements might also reveal information about the structure of the wind below the cloud level. These measurements are used to calculate

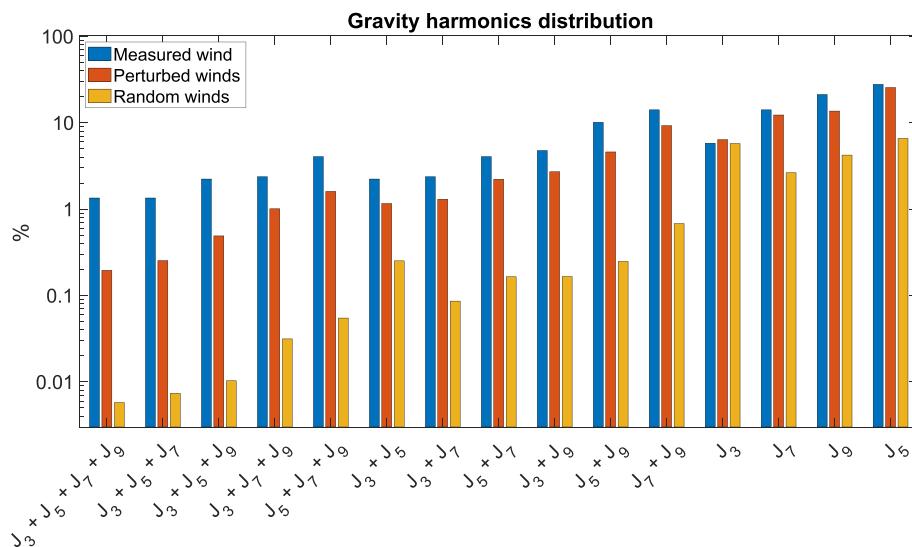


Figure 5. Summary of the solutions for the three presented cases of wind structures: Jupiter's measured wind at the cloud level (blue), 1,000 slightly modified meridional profiles of the zonal flow (red) and 1,000 random meridional profiles with a similar general structure to Jupiter's cloud-level profile (orange). The ordinate is a logarithmic scale of percentage relative to all the examined cases. The particular requirement of the solution to match the different odd gravity harmonics is presented by the abscissa.

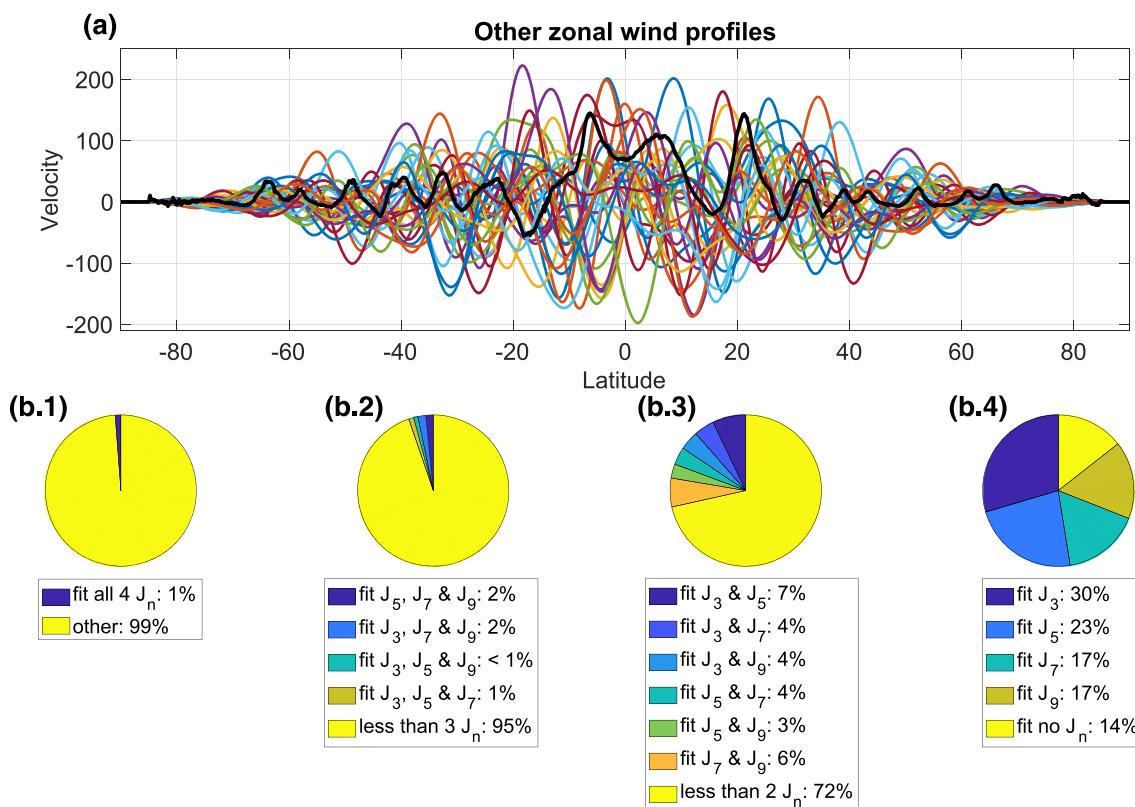


Figure 6. (a) Thirty examples of the 1,000 random zonal wind profiles examined (colors, $[m s^{-1}]$) and Jupiter's measured wind profile (black). (b.1–b.4) Summary of the random zonal flow meridional profiles' correspondence to the odd gravity harmonics. Only $\sim 1\%$ of the meridional profiles fit all four odd gravity harmonics (b.1), 5% of the profiles fit at least three odd gravity harmonics (b.2), 28% of the profiles fit at least two odd gravity harmonics (b.3), and 14% do not fit any of the odd gravity harmonics (b.4). The full compatibility distribution is detailed in the figure.

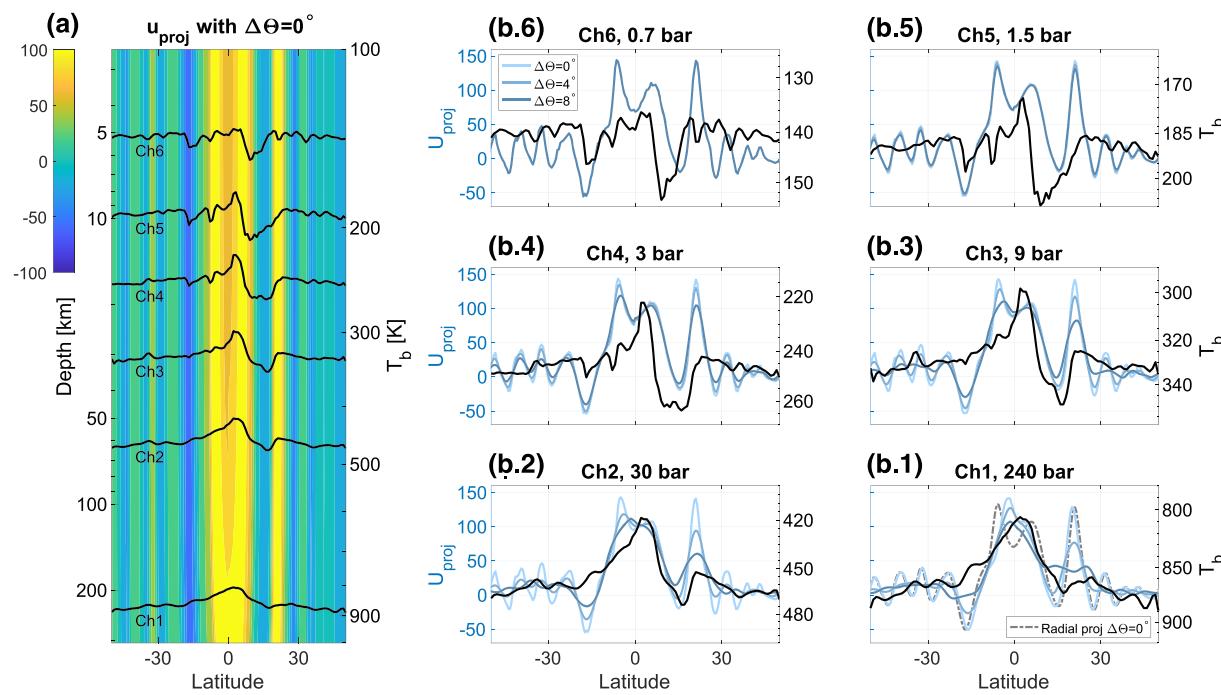


Figure 7. (a) Jupiter's projected wind velocities (m s^{-1}) between latitudes -50° and 50° in the upper 300 km of Jupiter (left ordinate) combined with the MWR nadir brightness temperature lines from Juno's PJ1 (black, right ordinate) in Channels 1 to 6, associated with frequencies of 0.6, 1.2, 2.6, 5.2, 10, and 22 GHz, respectively. (b.1–b.6) Jupiter's projected wind velocities (m s^{-1}) at Channel 1, 240 bar (b.1), Channel 2, 30 bar (b.2), Channel 3, 9 bar (b.3), Channel 4, 3 bar (b.4), Channel 5, 1.5 bar (b.5), and Channel 6, 0.6 bar (b.6) for a running average of $\Delta\theta=0^\circ$, $\Delta\theta=4^\circ$, and $\Delta\theta=8^\circ$ (darker blue with increasing $\Delta\theta$, left ordinate). Also shown is the brightness temperature [$^\circ\text{K}$] (black, right ordinate). The radial projection of the winds with no running average (dashed gray) is also presented in (b.1).

the nadir brightness temperature (T_b), a profile determined by the opacity of the atmosphere, which in Jupiter is determined mostly by ammonia abundance (Li et al., 2017). The MWR measurements reveal considerable variation of T_b with latitude and depth (Bolton et al., 2017) (Figure 7a, black lines). These variations with depth and the potential relation between T_b and the zonal jets (Ingersoll et al., 2017) suggest that Jupiter's zonal jets might be depth dependent, similarly to T_b , instead of simply projected inward (as in sections 3 and 4).

One approach for describing the relation between the brightness temperature and the zonal jets is taking the brightness temperature as simply temperature. Then, the relation is described by the thermal wind balance, as discussed in section 2. This approach, however, results in equatorial wind that is greater by 2 orders of magnitude than the measured cloud-level wind, which is unrealistic (Bolton et al., 2017). Another approach is taking the brightness temperature as an indicator for ammonia concentration (Ingersoll et al., 2017) and examining the relation to the zonal jets. As an example, such a relation is expected in the meridional circulation (Ferrel cells), where the cell-associated vertical velocity redistributes the substance and is accompanied by zonal jets (Fletcher et al., 2020). Here, we take the latter approach, analyzing a range of depth-dependent meridional profiles, compatible with the brightness temperature variations with depth.

Table 1

Correlation Coefficients Between ∇T_b and the Wind Velocity at Each Channel, for Winds With No Running Average at Depth ($\Delta\theta=0^\circ$), and Correlation Coefficients Between T_b and the Wind Velocity at the Same Depths With No Running Average ($\Delta\theta=0^\circ$), With Running Average of 4° ($\Delta\theta=4^\circ$), and With a Running Average of 8° ($\Delta\theta=8^\circ$)

Channel	u versus ∇T_b		u versus T_b	
	$\Delta\theta=0^\circ$	$\Delta\theta=4^\circ$	$\Delta\theta=0^\circ$	$\Delta\theta=8^\circ$
1	0.05	0.64	0.76	0.86
2	0.02	0.74	0.80	0.84
3	0.01	0.63	0.65	0.68
4	0.15	0.41	0.41	0.42
5	0.24	0.13	0.13	0.13
6	0.42	0.10	0.10	0.10

Note. For u versus T_b , the correlation increases with depth (or decreases with channel) and with running average.

When examining the correlation between T_b and the zonal jets, a different analysis should be taken at different latitudes and perhaps at different depths. If the zonal jets are associated with multiple Ferrel cells in alternating directions associated with regions of momentum convergence (eastward jets) and divergence (westward jets), a correlation is expected

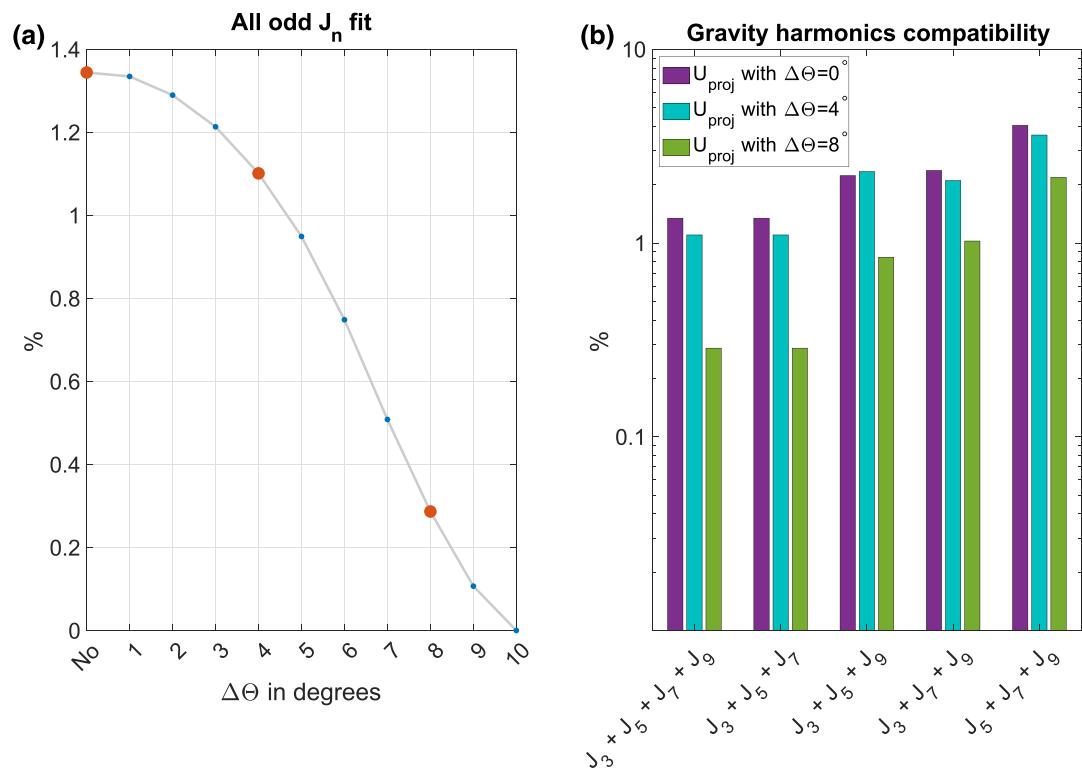


Figure 8. (a) The ability of the depth-dependent wind profiles to fit all four odd gravity harmonics (percentage of solutions) as a function of the smoothing factor in degrees (blue line). The three cases shown in Figure 7b are denoted by red dots. (b) The gravity harmonics distribution for the red dots is compatible with the three case studies in Figure 7b. The ordinate is a scale of percentage relative to the 5×10^5 decay options examined.

between the zonal velocity and the ammonia concentration gradient. In such a scenario, the meridional cells advect the ammonia concentration, maximizing its gradient where the jet peaks (Fletcher et al., 2020). However, momentum fluxes converging at the equator would lead to a superrotating jet (Kaspi et al., 2009) and might also lead to a maximal ammonia concentration. Therefore, at the equator, the zonal velocity is associated with the concentration itself and not with its gradient. These simple considerations motivates us to examine the correlation both between u and ∇T_b and between u and T_b (Table 1, Columns 2 and 3). Note that, as in all our experiments, the zonal jets are projected inward along the spin axis, as in section 3 (Figure 7a, colors). It is evident that the correlation between u and T_b is weak at the cloud level (Channel 6, 0.6 bar) but becomes stronger with depth (maximum at Channel 1, 240 bar), while the correlation between u and ∇T_b is strong at the cloud level and weakens with depth (at Channels 1–3, the correlation is weak). This alone might indicate on two opposite meridional cells, one stacked on top of the other (Fletcher et al., 2020; Showman & de Pater, 2005). At the cloud level, the correlation between u and ∇T_b improves if we do not consider the equatorial region, which is consistent with the Ferrel cells hypothesis. Projecting the winds in the radial direction instead of along cylinders does not improve the correlation to neither T_b or ∇T_b (Figure 7b.1).

The dominant feature leading to the strong correlation between u and T_b at Channel 1 is the equatorial anomaly, ascending at $\sim 15^\circ\text{S}$ and descending at $\sim 15^\circ\text{N}$ (Figure 7b.1). While at the cloud level, both the zonal jets and T_b reveal alternating patterns (Figure 7b.6), the waviness of T_b vanishes at deeper depths (Figures 7b.1 and 7b.2). Since T_b is depth dependent, getting smoother with depth from Channel 6 (cloud level) to Channel 1 (~ 300 km depth), we examine zonal jets that are depth dependent. Note that winds, projected along the spin axis, maintain their meridional profile with cylinders, and without further assumptions, are not depth dependent.

The modified (smoothed) wind at Channel 1 is composed using a running average of $\Delta\theta$ degrees latitude, where $\Delta\theta=0, 1, 2, \dots, 10^\circ$ (0° means that no running average is applied). The wind at Channel 6 is the

observed cloud-level profile; between Channel 1 and Channel 6, the wind strength is interpolated. Finally, the wind profile at the depth of 300 km (Channel 1) is projected inward along the spin axis with a decay profile as in the previous sections without further assumptions. In addition to the projected winds with no depth dependency ($\Delta\theta=0^\circ$, Figure 7b, light blue), we examine the correlation between u and T_b for two cases of a depth-dependent zonal wind ($\Delta\theta=4^\circ$ and $\Delta\theta=8^\circ$, Figure 7b, blue). Increasing the running average at depth improves the correlation at Channels 1–3 (Columns 3–5, Table 1), implying that the latitudinal variability of the jets might weaken beneath the cloud level.

Next, we examine the ability of the depth-dependent meridional profiles to explain the measured odd gravity harmonics. We examine a range of case studies, from slightly to largely modified depth-dependent profiles, until no solutions are found (Figure 8a). For slightly smoother profiles (small $\Delta\theta$), the ability to fit all four odd J_n is similar to that without any smoothing ($\Delta\theta=0^\circ$) (Figure 8a). Applying additional smoothing (increasing $\Delta\theta$) decreases the ability to fit the four odd J_n . Using more than a 10° running average results in no solutions for the odd gravity harmonics. The three case studies ($\Delta\theta=0^\circ$, $\Delta\theta=4^\circ$, and $\Delta\theta=8^\circ$) show a consistent trend when excluding one of the odd harmonics, such that the ability to fit the gravity measurements is reduced when it comes to smoother deep profiles (Figure 8b). This result is compatible with the previous case (section 4), indicating that deep wind that resembles the cloud-level wind can fit the gravity data, and changing the zonal wind structure considerably limits the ability to find a solution. The main conclusion of the analysis presented above is that wind profiles correlated with T_b at depth ($\Delta\theta=4^\circ$) can adequately fit the gravity measurements.

6. Discussion and Conclusions

The main challenge of interpreting the Juno gravity measurements is that the measurements provide only a handful of numbers (gravity harmonics), while the meridional and vertical profiles of the interior flow have many degrees of freedom. Therefore, by definition, the problem is ill posed. Acknowledging this inherent issue, Kaspi et al. (2018) used 4 degrees of freedom for the vertical flow profile (matching the number of the four odd harmonics) and found the best optimized profile for this allowed range. They addressed the nonuniqueness by showing the statistical likelihood of wind profiles for the interior that are completely different than the cloud-level flow. Kong et al. (2018) highlighted the nonuniqueness issue by showing that two different flow profiles can still satisfy the gravity measurements. In this study, we take a more methodological approach and consider a wider range of solutions and analyze their statistical likelihood. The flow profiles we consider, both for the meridional and vertical profiles of the zonal flow, are bound by physical considerations. We also address two main issues: First, all previous studies looked at all four odd gravity harmonics together and found the flow profiles best matching all four. Here, we investigate how each one of them separately bounds the flow. Second, in an attempt to coincide the gravity and microwave data, we explore whether deep profiles that are smoother than those of the cloud level, as possibility indicated by the Juno microwave measurements, can be consistent with the gravity measurements.

By assuming that the cloud-level wind profile is projected inward parallel to the spin axis, with some decay profile, we identify the envelope of possible solutions (Figure 1). We then relax the dependence on each of the odd gravity harmonics separately and analyze their individual contribution to the vertical profile of the zonal wind (Figure 2). We find that J_3 , the lowest-order odd harmonic that represents the dynamics of Jupiter, is sensitive at depths where the conductivity rises (beyond $\sim 3,000$ km), and the magnetic field might be interacting with the flow, resulting in the Lorentz force playing a key role in the dynamics. J_5 appears to be the most sensitive harmonic, giving a robust constraint on the vertical profile of the zonal flow alone (Figure 2b). Interestingly, J_9 does not give any new constraint on the flow if the other three harmonics are within the sensitivity range (Figure 2d). A possible explanation for the unique nature of J_5 comes from exploring the contribution function, which revealed that J_5 is most sensitive in the deeper regions, below 2,000 km (Figure 3).

The modified zonal flow analysis revealed a substantially bigger possible solution envelope than that obtained by extending the cloud-level wind (Figure 4). This implies that the zonal wind's structure may influence the depth sensitivity of each harmonic. Even for the perturbed winds, the flow cannot vanish at depths shallower than 2,000 km. The case with random winds implies that, with high probability, the wind cannot alter completely below the cloud level. Fitting the four odd gravity harmonics (or three if we

ignore J_9) requires either similar winds to the measured ones at the cloud level, that would penetrate a few thousand kilometers into the planet, or a very specific and statistically unlikely combination of meridional and decay profiles (Figures 5 and 6). Finally, the gravity harmonics induced by the slightly modified depth-dependent meridional profiles, which have a better correlation with the MWR measurements at depth (Figure 7 and Table 1), are still within Juno's gravity measurements' uncertainty, indicating that Jupiter's ammonia abundance could indeed reflect the profile of the zonal jet at 300 km (Figure 8).

Data Availability Statement

The Juno gravity measurements and MWR measurements are publicly available, see https://pds-atmospheres.nmsu.edu/data_and_services/atmospheres_data/JUNO/juno.html. Additional data can be found here <https://doi.org/10.5281/zenodo.3859828>.

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References

Bolton, S. J., Adriani, A., Adumitroaie, V., Allison, M., Anderson, J., Atreya, S., et al. (2017). Jupiter's interior and deep atmosphere: The initial pole-to-pole passes with the Juno spacecraft. *Science*, 356, 821–825. <https://doi.org/10.1126/science.aal2108>

Busse, F. H. (1976). A simple model of convection in the Jovian atmosphere. *Icarus*, 29, 255–260. [https://doi.org/10.1016/0019-1035\(76\)90053-1](https://doi.org/10.1016/0019-1035(76)90053-1)

Cao, H., & Stevenson, D. J. (2017). Zonal flow magnetic field interaction in the semi-conducting region of giant planets. *Icarus*, 296, 59–72. <https://doi.org/10.1016/j.icarus.2017.05.015>

Duer, K., Galanti, E., & Kaspi, Y. (2019). Analysis of Jupiter's deep jets combining Juno gravity and time-varying magnetic field measurements. *The Astrophysical Journal Letters*, 879(2), L22. <https://doi.org/10.3847/2041-8213/ab288e>

Fletcher, L. N., Kaspi, Y., Guillot, T., & Showman, A. P. (2020). How well do we understand the belt/zone circulation of giant planet atmospheres? *Space Science Reviews*, 216(2), 1–33. <https://doi.org/10.1007/s1214-019-0631-9>

French, M., Becker, A., Lorenzen, W., Nettelmann, N., Bethkenhagen, M., Wicht, J., & Redmer, R. (2012). Ab initio simulations for material properties along the Jupiter adiabat. *Astrophysical Journal Supplement*, 202, 5. <https://doi.org/10.1088/0067-0049/202/1/5>

Galanti, E., Cao, H., & Kaspi, Y. (2017). Constraining Jupiter's internal flows using Juno magnetic and gravity measurements. *Geophysical Research Letters*, 44, 8173–8181. <https://doi.org/10.1002/2017GL074903>

Galanti, E., Kaspi, Y., & Tziperman, E. (2017). A full, self-consistent, treatment of thermal wind balance on fluid planets. *Journal of Computational Physics*, 310, 175–195. <https://doi.org/10.1017/jfm.2016.687>

Garcia-Melendo, E., & Sánchez-Lavega, A. (2001). A study of the stability of Jovian zonal winds from HST images: 1995–2000. *Icarus*, 152(2), 316–330. <https://doi.org/10.1006/icar.2001.6646>

Gastine, T., Wicht, J., Duarte, L. D. V., Heimpel, M., & Becker, A. (2014). Explaining Jupiter's magnetic field and equatorial jet dynamics. *Geophysical Research Letters*, 41, 5410–5419. <https://doi.org/10.1002/2014GL060814>

Guillot, T., & Gautier, D. (2007). *Treatise of geophysics: 10. Planets and moons*. T. Spohn & J. Schubert (pp. 439–464). Elsevier.

Guillot, T., Miguel, Y., Militzer, B., Hubbard, W. B., Kaspi, Y., Galanti, E., et al. (2018). A suppression of differential rotation in Jupiter's deep interior. *Nature*, 555, 227–230. <https://doi.org/10.1038/nature25775>

Helled, R., Anderson, J. D., Podolak, M., & Schubert, G. (2010). Interior models of Uranus and Neptune. *The Astrophysical Journal*, 726(1), 15. <https://doi.org/10.1088/0004-637x/726/1/15>

Hubbard, W. B. (1984). *Planetary interiors* (pp. 343). New York: Van Nostrand Reinhold Co.

Hubbard, W. B. (1999). NOTE: Gravitational signature of Jupiter's deep zonal flows. *Icarus*, 137, 357–359. <https://doi.org/10.1006/icar.1998.6064>

Hubbard, W. B. (2012). High-precision Maclaurin-based models of rotating liquid planets. *The Astrophysical Journal Letters*, 756, L15. <https://doi.org/10.1088/2041-8205/756/1/L15>

Hubbard, W. B., Slattery, W. L., & Devito, C. L. (1975). High zonal harmonics of rapidly rotating planets. *The Astrophysical Journal*, 199, 504–516. <https://doi.org/10.1086/153717>

Hubbard, W. B., Trubitsyn, V. P., & Zharkov, V. N. (1974). Significance of gravitational moments for interior structure of Jupiter and Saturn. *Icarus*, 21(2), 147–151. [https://doi.org/10.1016/0019-1035\(74\)90130-4](https://doi.org/10.1016/0019-1035(74)90130-4)

Iess, L., Folkner, W. M., Durante, D., Parisi, M., Kaspi, Y., Galanti, E., et al. (2018). Measurement of Jupiter's asymmetric gravity field. *Nature*, 555(7695), 220–222. <https://doi.org/10.1038/nature25776>

Ingersoll, A. P., Adumitroaie, V., Allison, M. D., Atreya, S., Bellotti, A. A., Bolton, S. J., et al. (2017). Implications of the ammonia distribution on Jupiter from 1 to 100 bars as measured by the Juno microwave radiometer. *Geophysical Research Letters*, 44, 7676–7685. <https://doi.org/10.1002/2017GL074277>

Jones, C. A., & Kuzanyan, K. M. (2009). Compressible convection in the deep atmospheres of giant planets. *Icarus*, 204, 227–238. <https://doi.org/10.1016/j.icarus.2009.05.022>

Kaspi, Y. (2013). Inferring the depth of the zonal jets on Jupiter and Saturn from odd gravity harmonics. *Geophysical Research Letters*, 40, 676–680. <https://doi.org/10.1029/2012GL053873>

Kaspi, Y., Davighi, J. E., Galanti, E., & Hubbard, W. B. (2016). The gravitational signature of internal flows in giant planets: Comparing the thermal wind approach with barotropic potential-surface methods. *Icarus*, 276, 170–181. <https://doi.org/10.1016/j.icarus.2016.04.001>

Kaspi, Y., Flierl, G. R., & Showman, A. P. (2009). The deep wind structure of the giant planets: Results from an anelastic general circulation model. *Icarus*, 202, 525–542. <https://doi.org/10.1016/j.icarus.2009.03.026>

Kaspi, Y., Galanti, E., Hubbard, W. B., Stevenson, D. J., Bolton, S. J., Iess, L., et al. (2018). Jupiter's atmospheric jet-streams extend thousands of kilometers deep. *Nature*, 555, 223–226. <https://doi.org/10.1038/nature25793>

Kaspi, Y., Hubbard, W. B., Showman, A. P., & Flierl, G. R. (2010). Gravitational signature of Jupiter's internal dynamics. *Geophysical Research Letters*, 37, L01204. <https://doi.org/10.1029/2009GL041385>

Kong, D., Zhang, K., Schubert, G., & Anderson, J. D. (2018). Origin of Jupiter's cloud-level zonal winds remains a puzzle even after Juno. *Proceedings of the National Academy of Sciences of the United States of America*, 115(34), 8499–8504. <https://doi.org/10.1073/pnas.1805927115>

Li, C., Ingersoll, A., Janssen, M., Levin, S., Bolton, S., Adamitroaie, V., et al. (2017). The distribution of ammonia on Jupiter from a preliminary inversion of Juno microwave radiometer data. *Geophysical Research Letters*, 44, 5317–5325. <https://doi.org/10.1002/2017GL073159>

Liu, J., Goldreich, P. M., & Stevenson, D. J. (2008). Constraints on deep-seated zonal winds inside Jupiter and Saturn. *Icarus*, 196, 653–664. <https://doi.org/10.1016/j.icarus.2007.11.036>

Moore, K. M., Cao, H., Bloxham, J., Stevenson, D. J., Connerney, J. E. P., & Bolton, S. J. (2019). Time-variation of Jupiter's internal magnetic field consistent with zonal wind advection. *Nature Astronomy*, 3, 730–735. <https://doi.org/10.1038/s41550-019-0772-5>

Nellis, W. J., Mitchell, A. C., McCandless, P. C., Erskine, D. J., & Weir, S. T. (1992). Electronic energy gap of molecular hydrogen from electrical conductivity measurements at high shock pressures. *Physical Review Letters*, 68(19), 2937. <https://doi.org/10.1103/PhysRevLett.68.2937>

Nellis, W. J., Weir, S. T., & Mitchell, A. C. (1996). Metalization and electrical conductivity of hydrogen in Jupiter. *Science*, 273(5277), 936–938. <https://doi.org/10.1126/science.273.5277.936>

Nettelmann, N., Helled, R., Fortney, J. J., & Redmer, R. (2013). New indication for a dichotomy in the interior structure of Uranus and Neptune from the application of modified shape and rotation data. *Planetary and Space Science*, 77, 143–151. <https://doi.org/10.1016/j.pss.2012.06.019>

Nettelmann, N., Piustow, R., & Redmer, R. (2013). Saturn layered structure and homogeneous evolution models with different EOSs. *Icarus*, 225(1), 548–557. <https://doi.org/10.1016/j.icarus.2013.04.018>

Pedlosky, J. (1987). *Geophysical fluid dynamics* (pp. 710): Springer-Verlag.

Salyk, C., Ingersoll, A. P., Lorre, J., Vasavada, A., & Del Genio, A. D. (2006). Interaction between eddies and mean flow in Jupiter's atmosphere: Analysis of Cassini imaging data. *Icarus*, 185, 430–442. <https://doi.org/10.1016/j.icarus.2006.08.007>

Showman, A. P., & de Pater, I. (2005). Dynamical implications of Jupiter's tropospheric ammonia abundance. *Icarus*, 174, 192–204. <https://doi.org/10.1016/j.icarus.2004.10.004>

Showman, A. P., & Kaspi, Y. (2013). Atmospheric dynamics of brown dwarfs and directly imaged giant planets. *The Astrophysical Journal*, 776, 85–103. <https://doi.org/10.1088/0004-637x/776/2/85>

Tollefson, J., Wong, M. H., de Pater, I., Simon, A. A., Orton, G. S., Rogers, J. H., et al. (2017). Changes in Jupiter's zonal wind profile preceding and during the Juno mission. *Icarus*, 296, 163–178. <https://doi.org/10.1016/j.icarus.2017.06.007>

Weir, S. T., Mitchell, A. C., & Nellis, W. J. (1996). Metalization of fluid molecular hydrogen at 140 GPa (1.4 mbar). *Physical Review Letters*, 76, 1860–1863. <https://doi.org/10.1103/PhysRevLett.76.1860>

Wicht, J., Gastine, T., & Duarte, L. D. (2019). Dynamo action in the steeply decaying conductivity region of Jupiter-like dynamo models. *Journal of Geophysical Research: Planets*, 124, 837–863. <https://doi.org/10.1029/2018JE005759>

Wicht, J., Gastine, T., Duarte, L. D., & Dietrich, W. (2019). Dynamo action of the zonal winds in Jupiter. *Astronomy & Astrophysics*, 629, A125. <https://doi.org/10.1051/0004-6361/201935682>

Zharkov, V. N., & Trubitsyn, V. P. (1974). Determination of the equation of state of the molecular envelopes of Jupiter and Saturn from their gravitational moments. *Icarus*, 21(2), 152–156. [https://doi.org/10.1016/0019-1035\(74\)90131-6](https://doi.org/10.1016/0019-1035(74)90131-6)

Kaspi, Y., Galanti, E., Showman, A. P., Stevenson, D. J., Guillot, T., Iess, L., & Bolton, S. J. (2020). Comparison of the deep atmospheric dynamics of Jupiter and Saturn in light of the Juno and Cassini gravity measurements. *Space Science Reviews*, 216(84). <https://doi.org/10.1007/s11214-020-00705-7>